# What was the Universe before the hot Big Bang?

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- With Big Bang nucleosynthesis theory and observations

  we are confident of the theory of the early Universe

  at temperatures up to $T \simeq 1$ MeV, age $t \simeq 1$ second

- With the LHC and dark matter detection,

  we hope to be able to go

  up to temperatures $T \sim 100$ GeV, age $t \sim 10^{-10}$ second

- Extrapolation back in time with known laws of physics and known elementary particles and fields $\Longrightarrow$ hot Universe, starts from Big Bang singularity (infinite temperature, infinite expansion rate)

We now KNOW that this is not the whole story.

# Key: cosmological perturbations

- Our Universe is not exactly homogeneous.

Inhomogeneities: ⊙ density perturbations and associated gravitational potentials (3d scalar), observed;

⊙ gravitational waves (3d tensor), not observed (yet?).

Today: inhomogeneities strong and non-linear

In the past: amplitudes small,

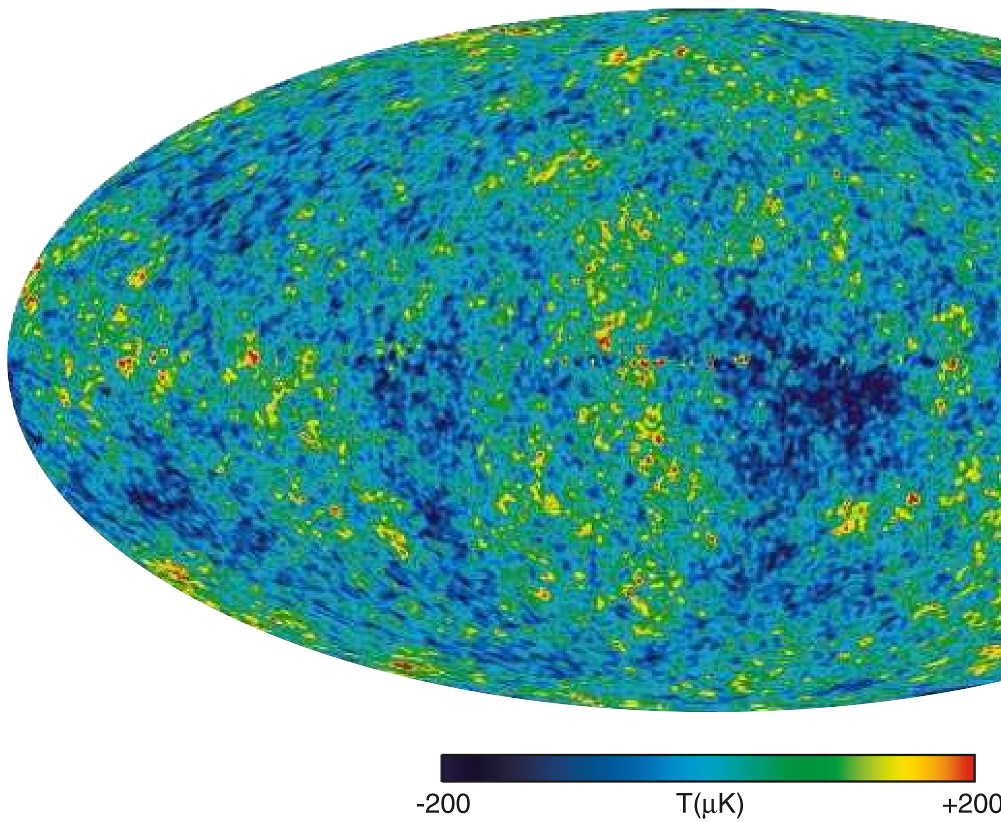
$$\frac{\delta\rho}{\rho} = 10^{-4} - 10^{-5}$$

Linear analysis appropriate.

## How are they measured?

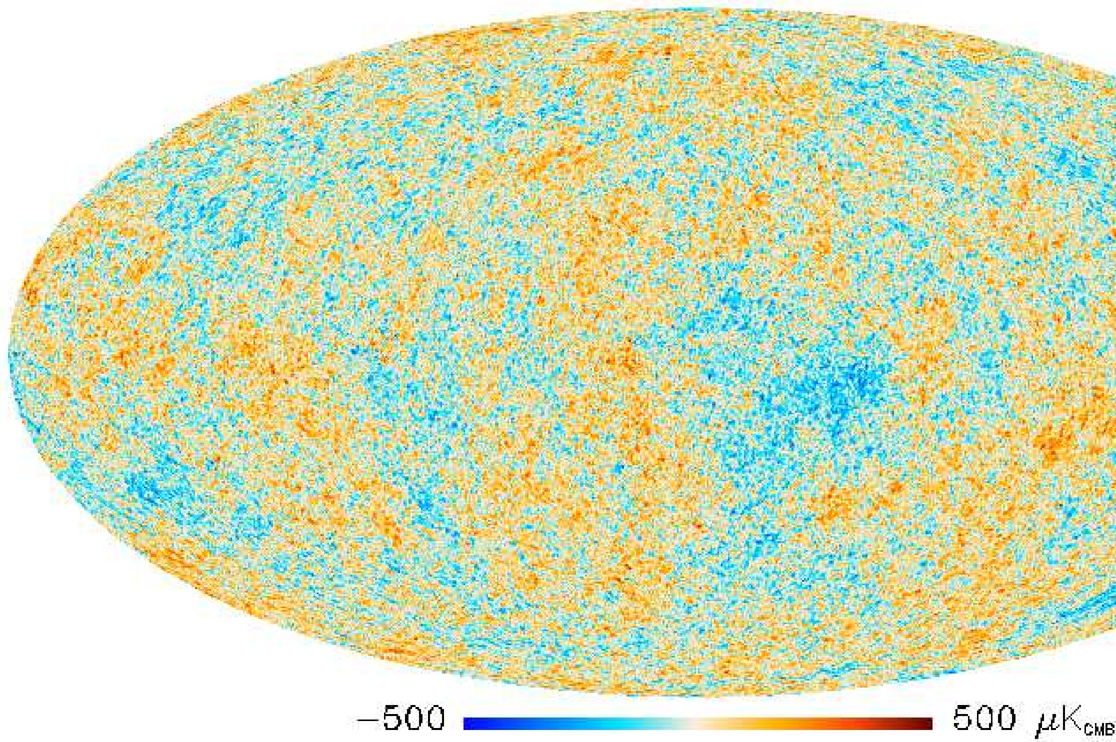
- Cosmic microwave background: photographic picture of the Universe at age 380 000 yrs, $T = 3000$ K (transition from plasma to neutral gas, mostly hydrogen and helium)
  - Temperature anisotropy
  - Polarization
- Deep surveys of galaxies and quasars, cover good part of entire visible Universe
- Gravitational lensing, etc.

$$T = 2.726^\circ K, \quad \frac{\delta T}{T} \sim 10^{-4} - 10^{-5}$$



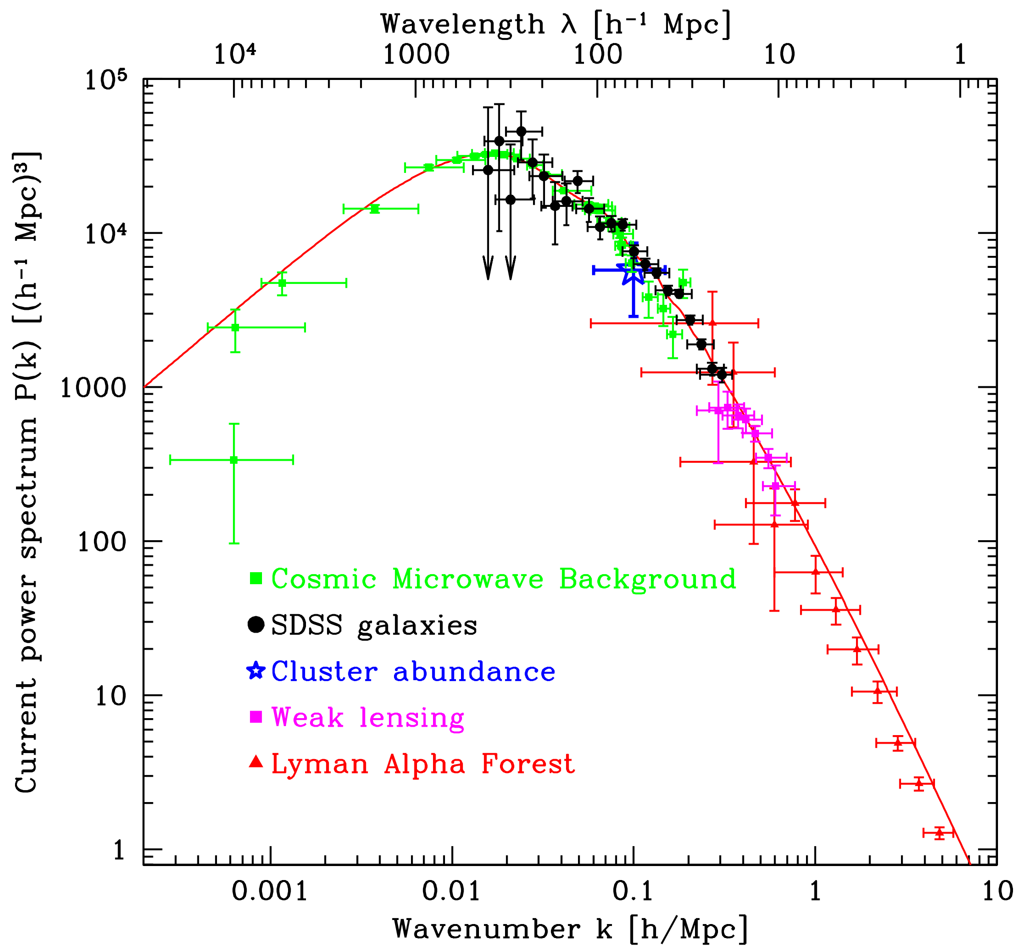

WMAP

$$T = 2.726^{\circ}K, \quad \frac{\delta T}{T} \sim 10^{-4} - 10^{-5}$$




Planck

# Overall consistency




- NB: density perturbations = random field.
  $k$ = wavenumber
  $P(k)$ = power spectrum transfered to present epoch using linear theory

- Properties of perturbations in conventional hot Universe.

Friedmann–Lemaître–Robertson–Walker metric:

$$ds^2 = dt^2 - a^2(t) d\vec{x}^{\,2}$$

Expanding Universe:

$a(t) \propto t^{1/2}$ at radiation domination stage (before $T \simeq 1$ eV, $t \simeq 60$ thousand years)

$a(t) \propto t^{2/3}$ at matter domination stage (until recently).

Cosmological horizon (assuming that nothing preceeded hot epoch): length that light travels from Big Bang moment,

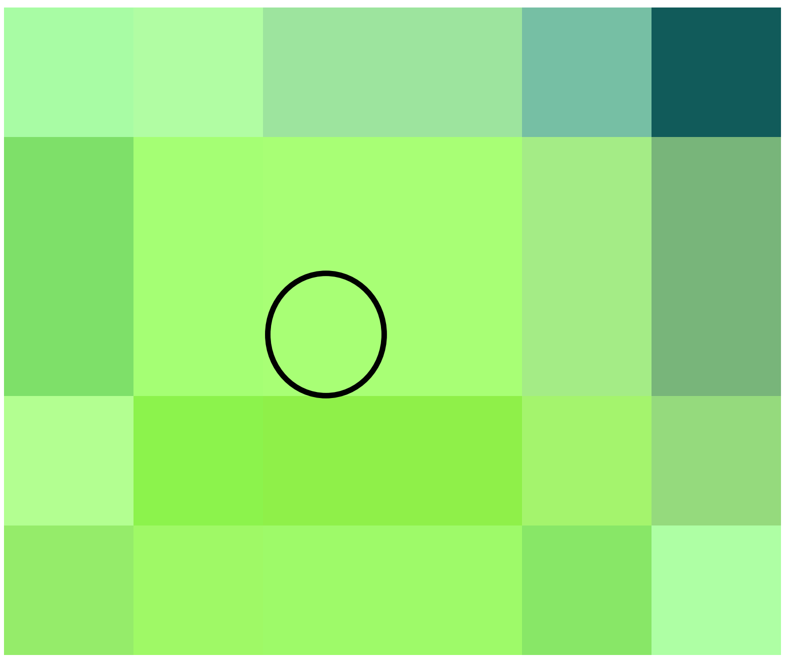
$$l_H(t) = (2-3)t$$

- Wavelength of perturbation grows as $a(t)$.
  E.g., at radiation domination

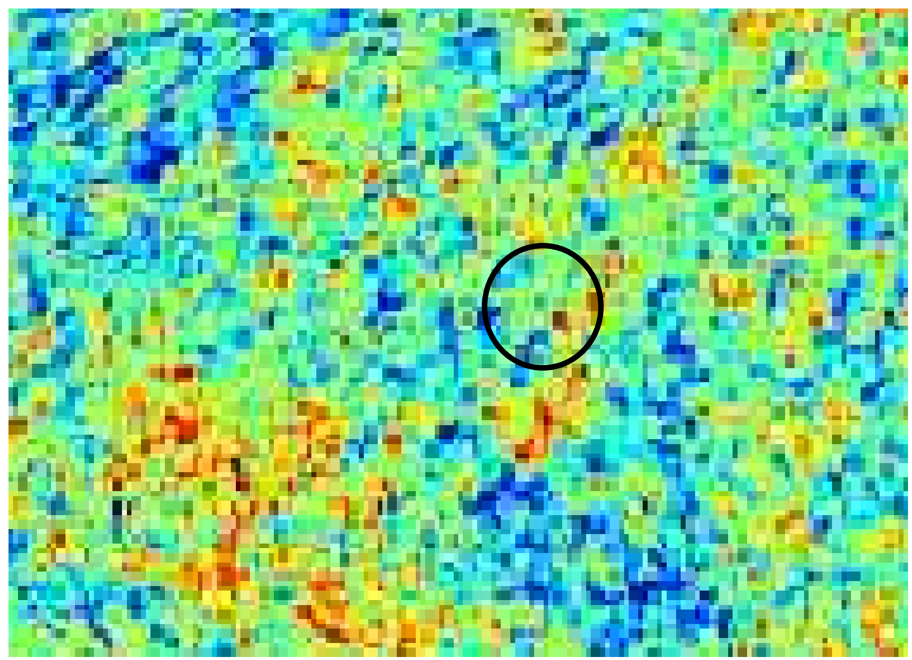
$$\lambda(t) \propto t^{1/2} \quad \text{while} \quad l_H \propto t$$

Today $\lambda < l_H$, subhorizon regime

Early on $\lambda(t) > l_H$, superhorizon regime.



superhorizon mode



subhorizon mode

- In other words, physical wavenumber (momentum) gets redshifted,

$$q(t) = \frac{2\pi}{\lambda(t)} = \frac{k}{a(t)}\,, \qquad k = \text{const} = \text{coordinate momentum}$$

Today

$$q > H \equiv \frac{\dot{a}}{a}$$

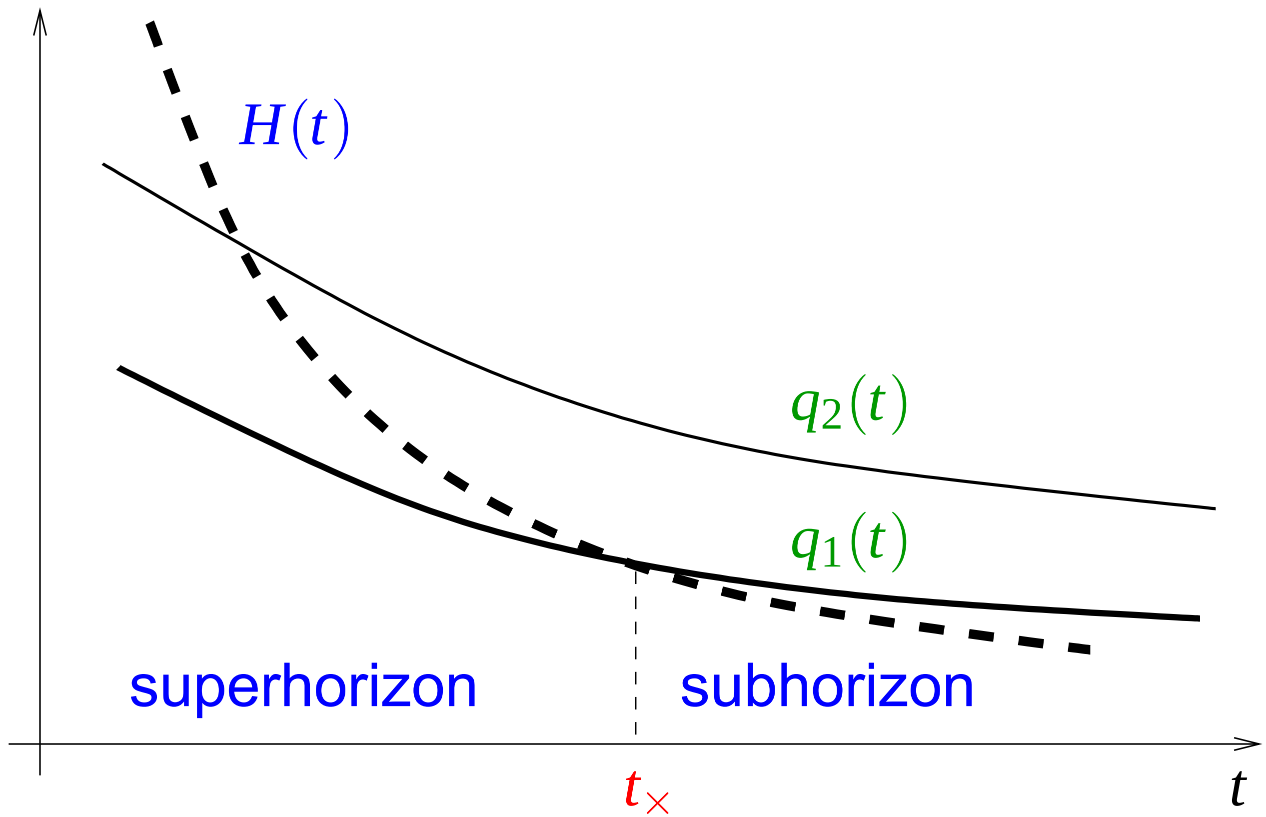
Early on

$$q(t) < H(t)$$

Very different regimes of evolution.

NB: Horizon entry occured after Big Bang Nucleosynthesis epoch for modes of all relevant wavelengths $\Longleftrightarrow$ no guesswork at this point.

# Regimes at radiation (and matter) domination




$q_2 > q_1$

# Major issue: origin of perturbations

- Causality $\Longrightarrow$ perturbations can be generated only when they are subhorizon.

Off-hand possibilities:

- Perturbations were never superhorizon, they were generated at the hot cosmological epoch by some causal mechanism.

  E.g., seeded by topological defects (cosmic strings, etc.)

  N. Turok et.al.' 90s

  The only possibility, if expansion started from hot Big Bang.

  No longer an option!

- Hot epoch was preceeded by some other epoch. Perturbations were generated then.

- Perturbations in baryon-photon plasma = sound waves.

  If they were superhorizon, they started off with one and the same phase.

  Reason: solutions to wave equation in superhorizon regime in expanding Universe

$$\frac{\delta\rho}{\rho} = \text{const} \quad \text{and} \quad \frac{\delta\rho}{\rho} = \frac{\text{const}}{t^{3/2}}$$

  Assume that modes were superhorizon. If the Universe was not very inhomogeneous at early times, the initial condition is unique (up to amplitude),

$$\frac{\delta\rho}{\rho} = \text{const} \implies \frac{d}{dt}\frac{\delta\rho}{\rho} = 0$$

  Acoustic oscillations start after entering the horizon at zero velocity of medium $\implies$ phase of oscillations uniquely defined.

- Perturbations develop different phases by the time of photon last scattering ( = recombination), depending on wave vector:

$$\frac{\delta\rho}{\rho}(t_r) \propto \cos\left(\int_0^{t_r} dt\ v_s\ q(t)\right)$$

($v_s$ = sound speed in baryon-photon plasma) $\Longrightarrow$

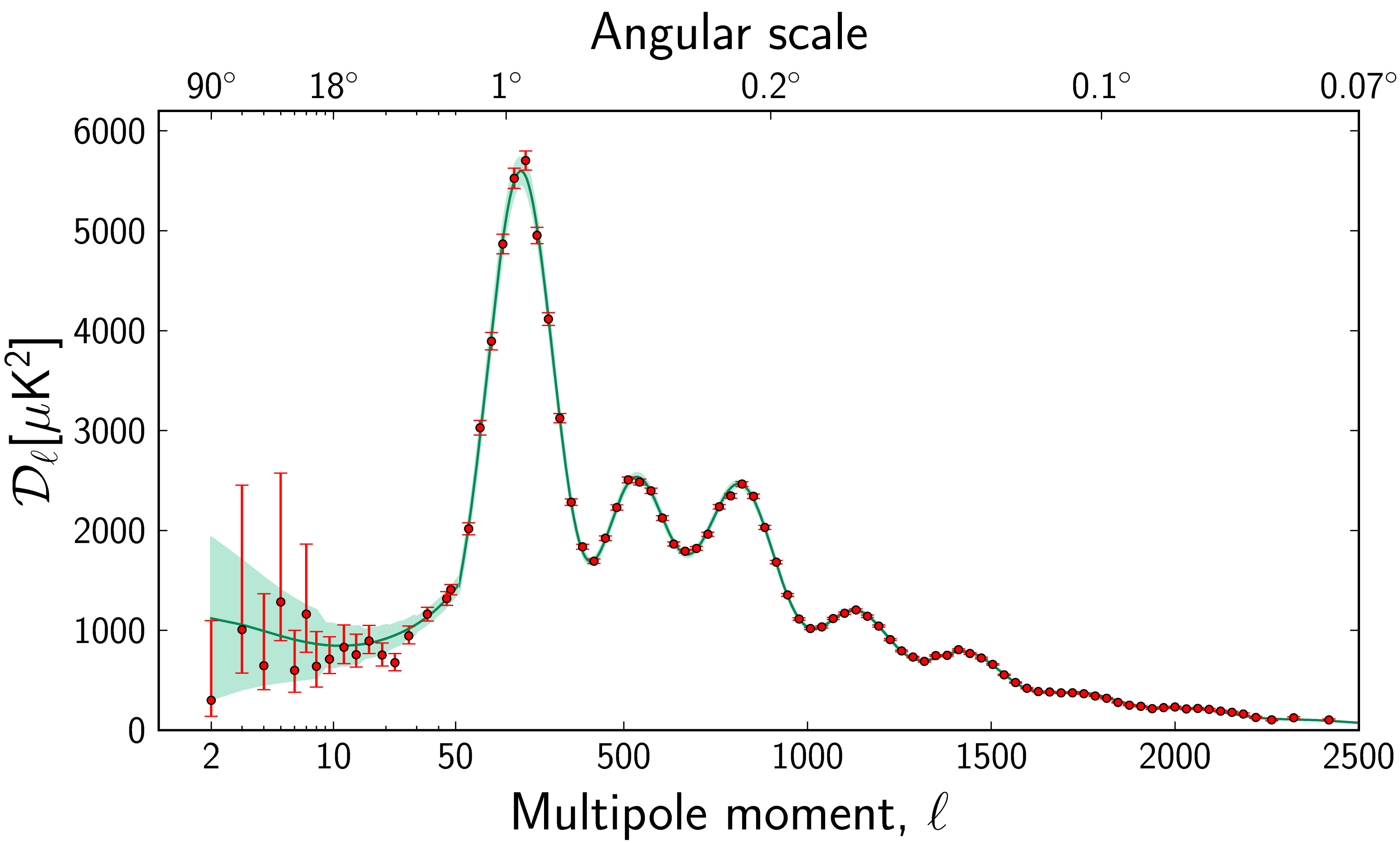
Oscillations in CMB temperature angular spectrum

Fourier decomposition of temperatue fluctuations:

$$\delta T(\theta,\varphi) = \sum_{l,m} a_{lm} Y_{lm}(\theta,\varphi)$$

$\langle a^*_{lm} a_{lm}\rangle = C_l$, temperature angular spectrum; $D_l = l(l+1)C_l$

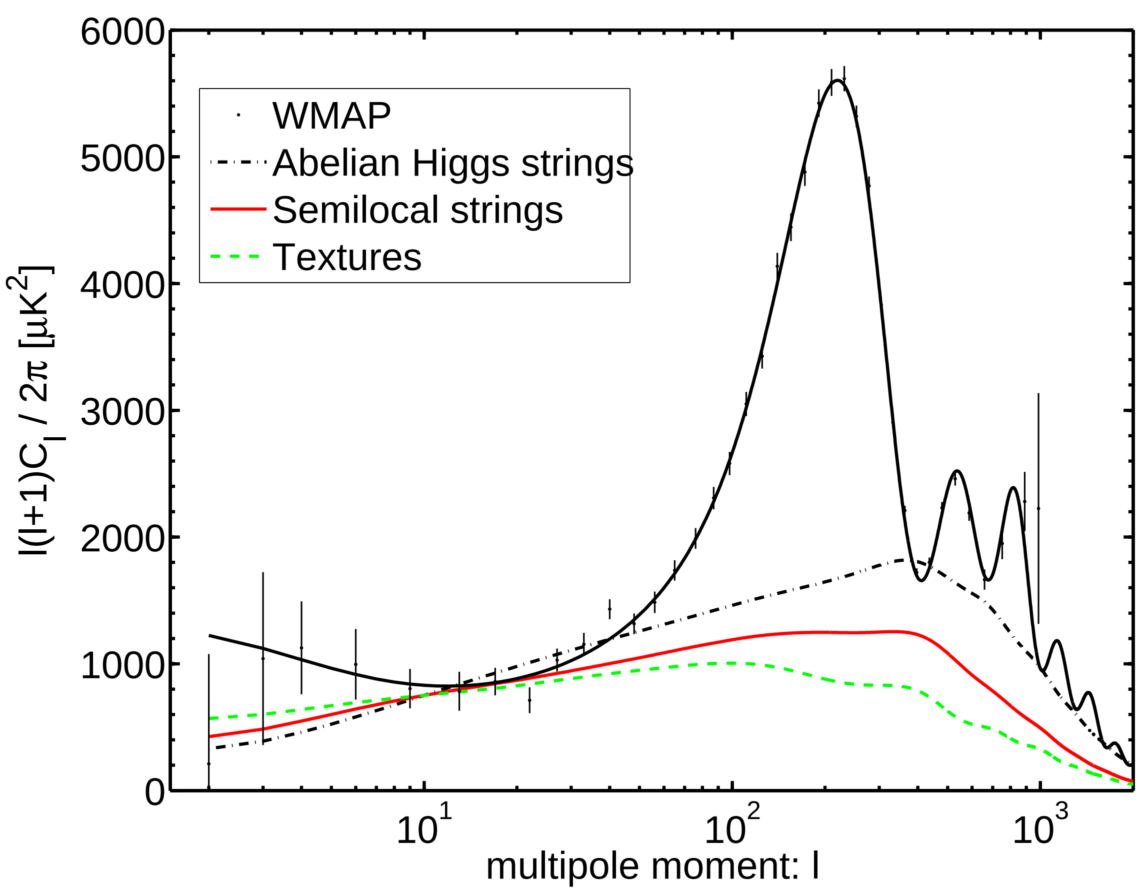
larger $l \Longleftrightarrow$ smaller angular scales, shorter wavelengths



Angular scale
90°
18°
1°
0.2°
0.1°
0.07°
$\mathcal{D}_\ell [\mu K^2]$
6000
5000
4000
3000
2000
1000
0
2
10
50
500
1000
1500
2000
2500
Multipole moment, $\ell$

- Furthermore, there are perturbations which were superhorizon at the time of photon last scattering

These properties would not be present if perturbations were generated at hot epoch in causal manner.



Primordial perturbations were generated at some yet unknown epoch before the hot expansion stage.

That epoch must have been long and unusual:
perturbations were subhorizon early at that epoch,
visible part of the Universe was in a causally connected region.

- Excellent guess: inflation

Starobinsky’79; Guth’81; Linde’82; Albrecht and Steinhardt’82

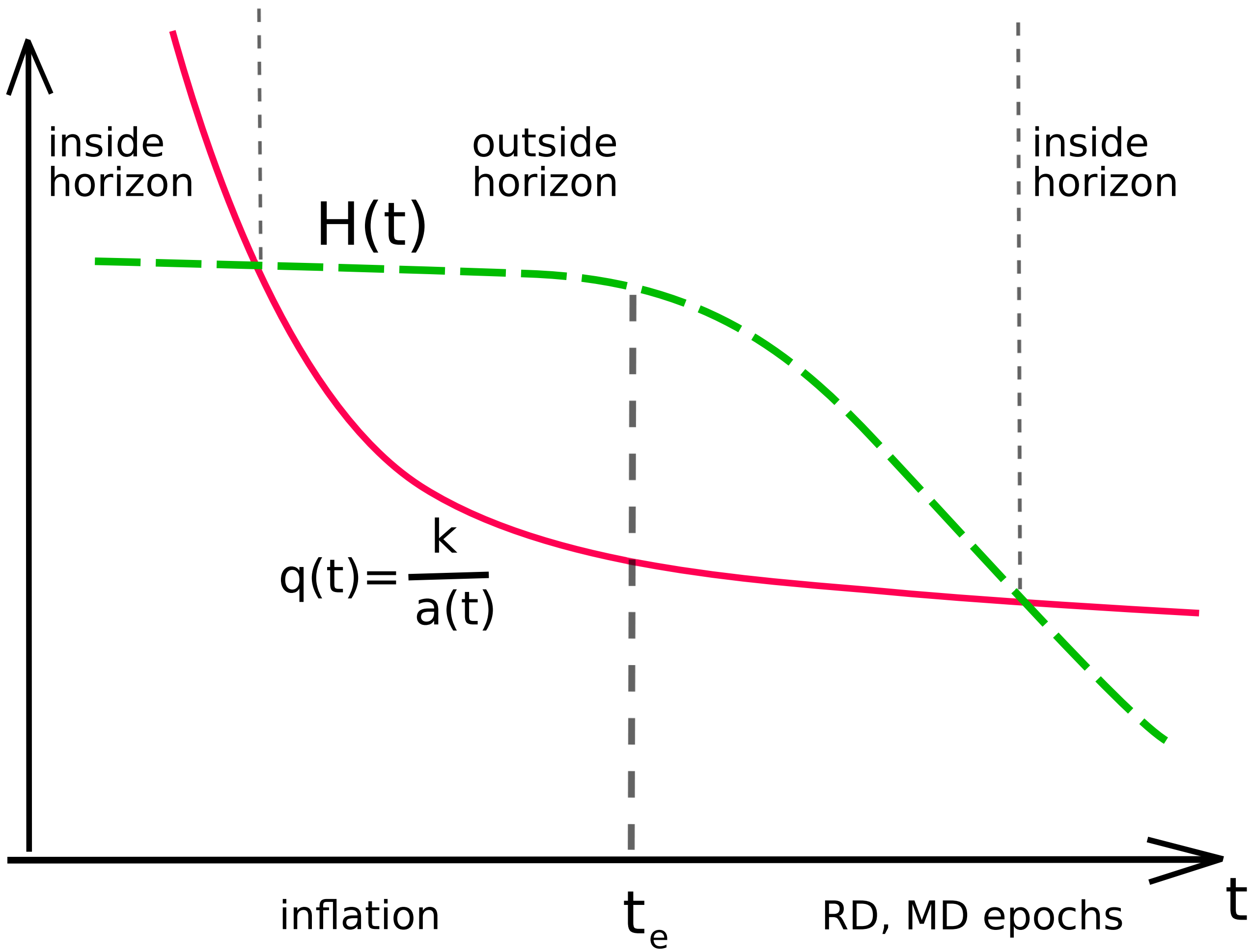
Exponential expansion with almost constant Hubble rate,

$$a(t) = \mathrm{e}^{\int H dt}\,, \qquad H \approx \mathrm{const}$$

Perturbations subhorizon early at inflation:

$$q(t) = \frac{k}{a(t)} \gg H$$

## Physical wave number and Hubble parameter at inflation and later:



# Alternatives to inflation:

- Contraction — Bounce — Expansion
- Start up from static state

Creminelli et.al.'06; '10

Difficult, though not impossible. Einstein equations (neglecting spatial curvature)

$$H^2 = \frac{8\pi}{3} G\rho$$
$$\frac{dH}{dt} = -4\pi(\rho + p)$$

$\rho =$ energy density, $p =$ pressure, $H = \dot{a}/a$.

Bounce, start up scenarios $\Longrightarrow \frac{dH}{dt} > 0 \Longrightarrow \rho > 0$ and $p < -\rho$

Exotic matter. Potential problems with instabilities, superluminal propagation/causality. Solvable.

## Other observational facts about density perturbations

(valid within certain error bars!)

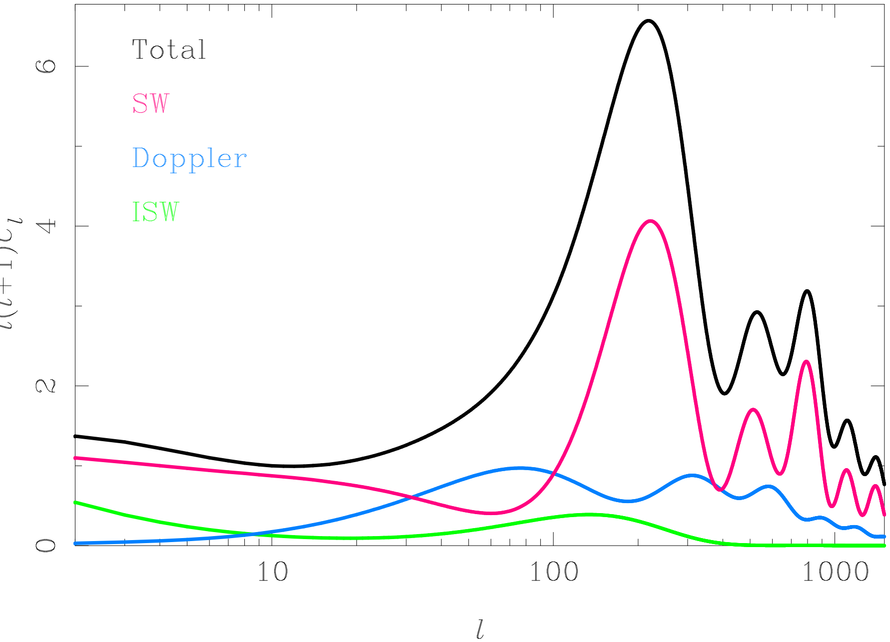
- Perturbations in overall density, not in composition (jargon: “adiabatic”)

$$\frac{\text{baryon density}}{\text{entropy density}} = \frac{\text{dark matter density}}{\text{entropy density}} = \text{const in space}$$

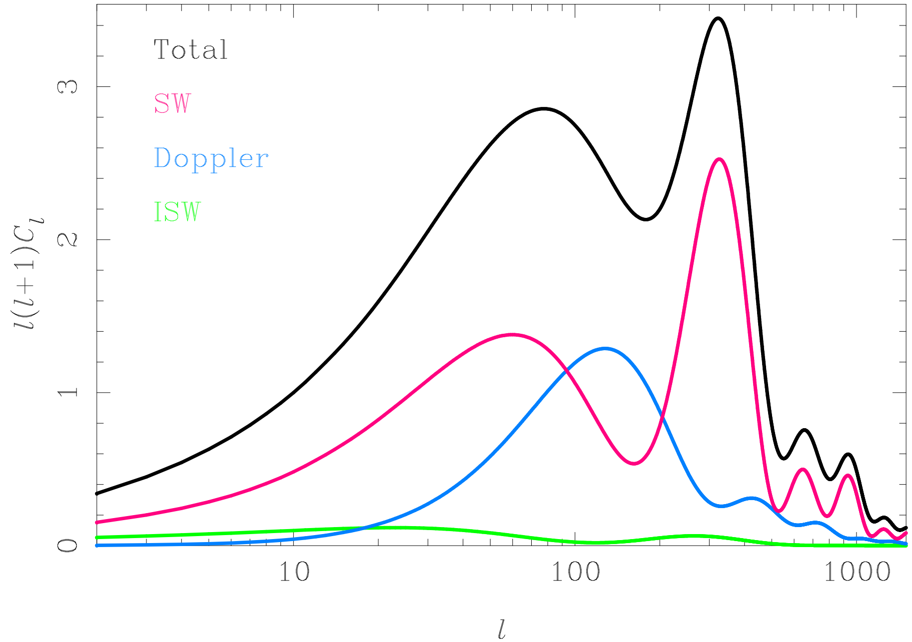
Consistent with generation of baryon asymmetry and dark matter at hot stage.

Perturbation in chemical composition (jargon: “isocurvature” or “entropy”) $\Longrightarrow$ wrong initial condition for acoustic oscillations $\Longrightarrow$ wrong prediction for CMB angular spectrum.

# CMB angular spectra



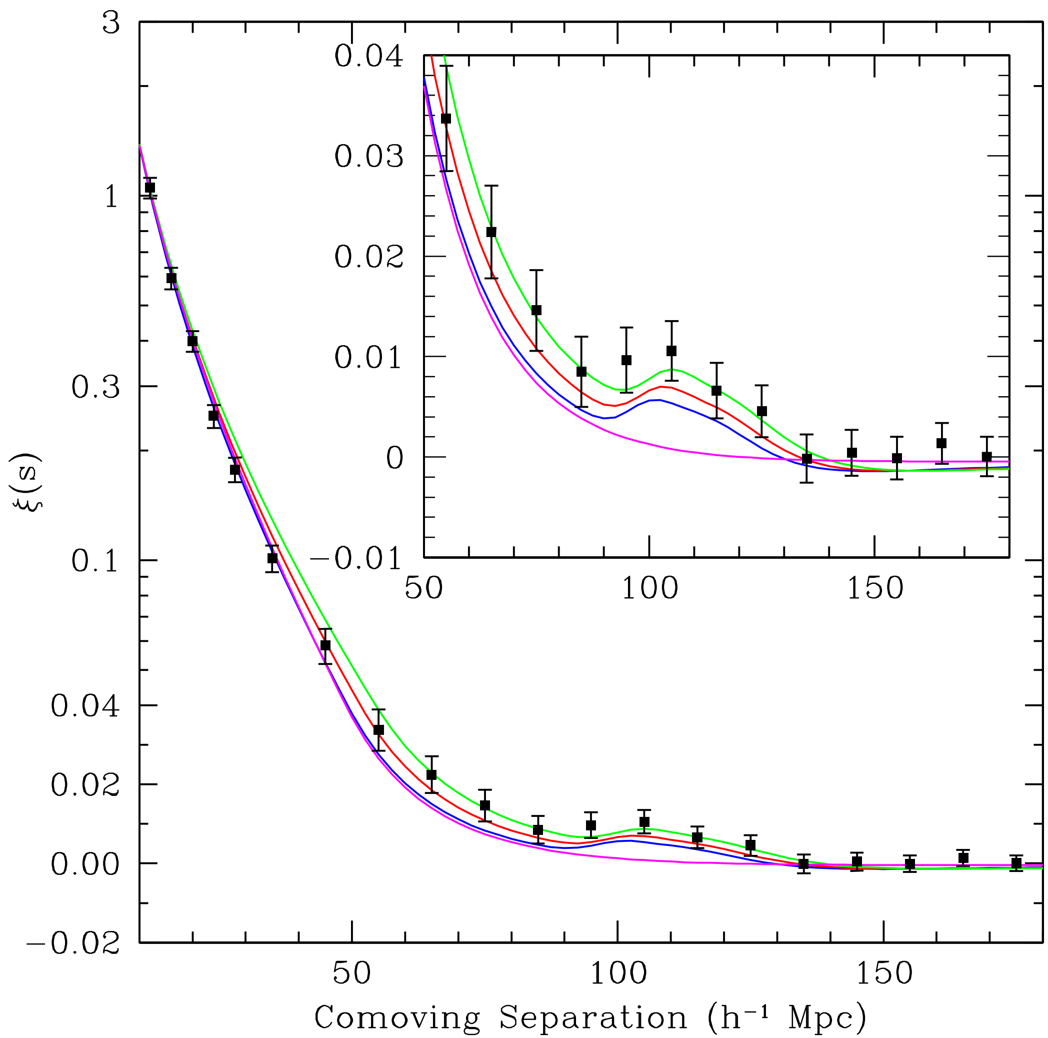

homogeneous chemical composition




inhomogeneous chemical composition

- NB: even weak variation of composition over space would mean exotic mechanism of baryon asymmetry and/or dark matter generation

# Baryon Acoustic Oscillations in galaxy correlation function



- Primordial perturbations are Gaussian.

  Gaussianity = Wick theorem for correlation functions

  This suggests the origin: enhanced vacuum fluctuations of weakly coupled quatum field(s)

  NB: Linear evolution does not spoil Gaussianity.

  - Inflation does the job very well: fluctuations of all light fields get enhanced greatly due to fast expansion of the Universe.

    Including the field that dominates energy density (inflaton) $\Longrightarrow$ perturbations in energy density.

    Mukhanov, Chibisov'81; Hawking'82; Starobinsky'82; Guth, Pi'82; Bardeen et.al.'83

  - Enhancement of vacuum fluctuations is less automatic in alternative scenarios

## Primordial power spectrum is almost flat.

Homogeneity and anisotropy of Gaussian random field:

$$\langle \frac{\delta\rho}{\rho}(\vec{k}) \frac{\delta\rho}{\rho}(\vec{k}') \rangle = \frac{1}{4\pi k^3} \mathscr{P}(k) \delta(\vec{k}+\vec{k}')$$

$\mathscr{P}(k) =$ power spectrum, gives fluctuation in logarithmic interval of momenta,

$$\langle \left( \frac{\delta\rho}{\rho}(\vec{x}) \right)^2 \rangle = \int_0^\infty \frac{dk}{k} \mathscr{P}(k)$$

Flat spectrum: $\mathscr{P}$ is independent of $k$

Harrison' 70; Zeldovich' 72

Parametrization

$$\mathscr{P}(k) = A \left( \frac{k}{k_*} \right)^{n_s - 1}$$

$A =$ amplitude, $(n_s - 1) =$ tilt, $k_* =$ fiducial momentum (matter of convention). Flat spectrum: $n_s = 1$. Exp.: $n_s = 0.96$.

## There must be some symmetry behind flatness of spectrum

- Inflation: symmetry of de Sitter space-time

$$ds^2 = dt^2 - \mathrm{e}^{2Ht} d\vec{x}^{\,2}$$

Symmetry: spatial dilatations supplemented by time translations

$$\vec{x} \to \lambda \vec{x}\,, \quad t \to t - \frac{1}{2H} \log \lambda$$

Inflation automatically generates nearly flat spectrum.

- Alternative: conformal symmetry

Conformal group includes dilatations, $x^\mu \to \lambda x^\mu$.

$\Longrightarrow$ No scale, good chance for flatness of spectrum

V.R.' 09;

Creminelli, Nicolis, Trincherini' 10

- NB: Conformal symmetry has long been discussed in the context of Quantum Field Theory and particle physics.

  Particularly important in the context of supersymmetry: many interesting superconformal theories.

  Large and powerful symmetry behind, e.g., adS/CFT correspondence and a number of other QFT phenomena

  Maldacena' 97

  It may well be that ultimate theory of Nature is (super)conformal

  What if our Universe started off from a conformal state and then evolved to much less symmetric state we see today?

  Exploratory stage: toy models so far.

# Can one tell?

More intricate properties of cosmological perturbations

Not detected yet.

- Primordial gravitational waves

Sizeable amplitude, (almost) flat power spectrum predicted by simplest (and hence most plausible) inflationary models
but not alternatives to inflation

May make detectable imprint on CMB temperature anisotropy

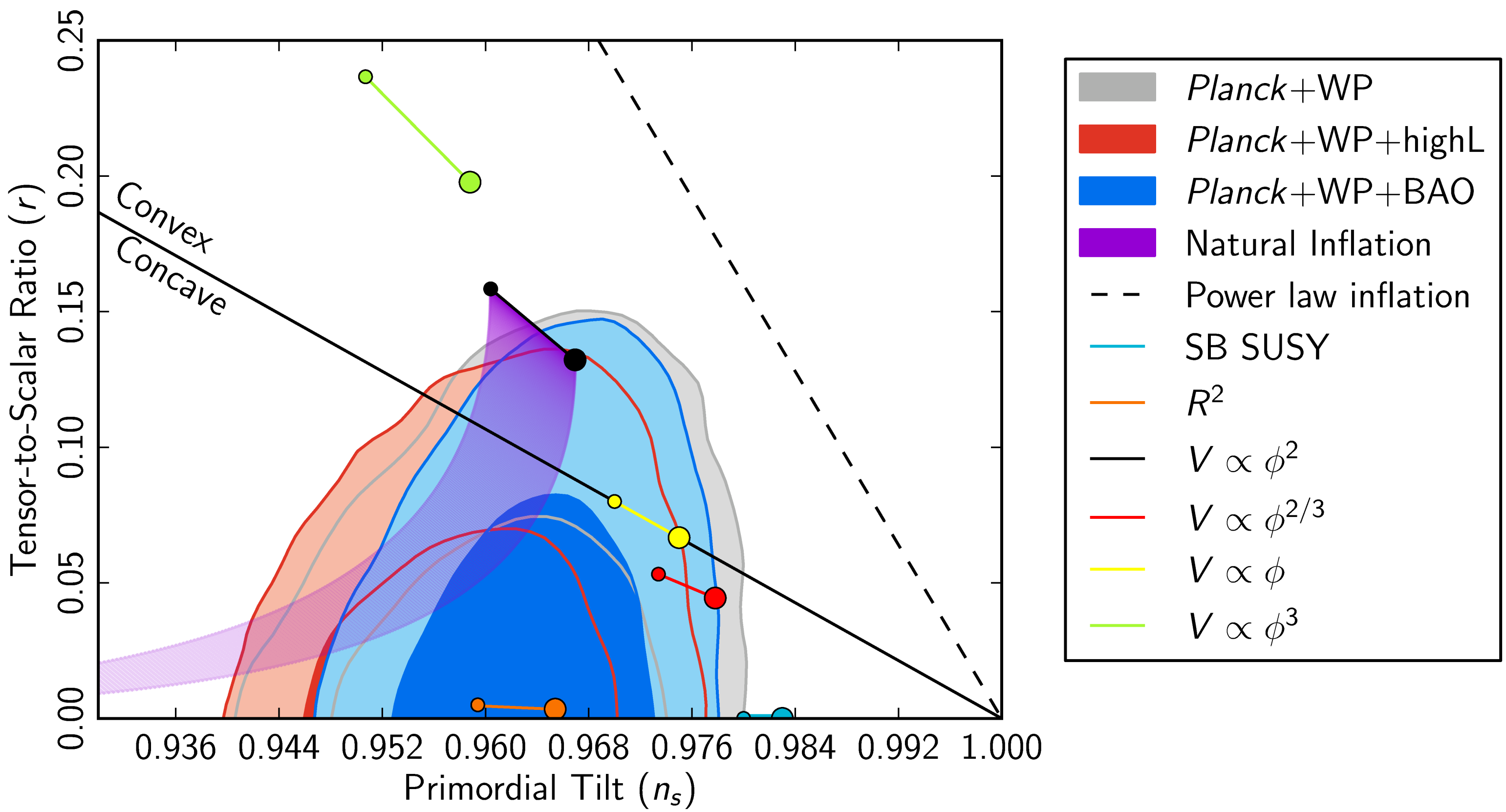
V.R., Sazhin, Veryaskin' 82;
Fabbri, Pollock' 83; ...

and especially on CMB polarization

Kamionkowski, Kosowsky, Stebbins' 96;
Seljak, Zaldarriaga' 96; ...

Smoking gun for inflation

# Scalar tilt vs tensor power




NB:

$$r = \left( \frac{\text{amplitude of gravity waves}}{\text{amplitude of density perturbations}} \right)^2$$

- Non-Gaussianity
  - Very small in the simplest inflationary theories
  - Sizeable in more contrived inflationary models and in alternatives to inflation. Often begins with bispectrum

$$\langle \frac{\delta\rho}{\rho}(\mathbf{k_1}) \frac{\delta\rho}{\rho}(\mathbf{k_2}) \frac{\delta\rho}{\rho}(\mathbf{k_3}) \rangle = \delta(\mathbf{k_1} + \mathbf{k_2} + \mathbf{k_3}) G(k_i^2, \mathbf{k_1 k_2}, \mathbf{k_1 k_3})$$

    Shape of $G(k_i^2, \mathbf{k_1 k_2}, \mathbf{k_1 k_3})$ different in different models $\Longrightarrow$ potential discriminator.
  - Sometimes bispectrum vanishes, e.g., due to some symmetries: $\phi \to \phi^*$ in conformal model. But trispectrum (connected 4-point function) may be measurable.

- Statistical anisotropy

$$\mathscr{P}(\mathbf{k}) = \mathscr{P}_0(k)\left(1 + w_{ij}(k)\frac{k_i k_j}{k^2} + \dots\right)$$

  - Anisotropy of the Universe at pre-hot stage
  - Possible in inflation with strong vector fields (rather contrived). Quadrupole only.

    Ackerman, Carroll, Wise' 07; Pullen, Kamionkowski' 07;
    Watanabe, Kanno, Soda' 09

  - Natural in some other scenarios, including conformal model. May begin with dipole.

    Libanov, V.R.' 10; Libanov, Ramazanov, V.R., in progress

  - Would show up in correlators

$$\langle a_{lm} a_{l'm'} \rangle \quad \text{with} \;\; l' \neq l \;\; \text{and/or} \;\; m' \neq m$$

# To summarize:

- Available data on cosmological perturbations (notably, CMB anisotropies) give confidence that the hot stage of the cosmological evolution was preceeded by some other epoch, at which these perturbations were generated.

- Inflation is consistent with all data. But there are competitors: the data may rather be viewed as pointing towards early conformal epoch of the cosmological evolution.

More options:

Matter bounce, Finelli, Brandenberger' 01.

Negative exponential potential, Lehners et. al.' 07;
Buchbinder, Khouri, Ovrut' 07; Creminelli, Senatore' 07.

Lifshitz scalar, Mukohyama' 09

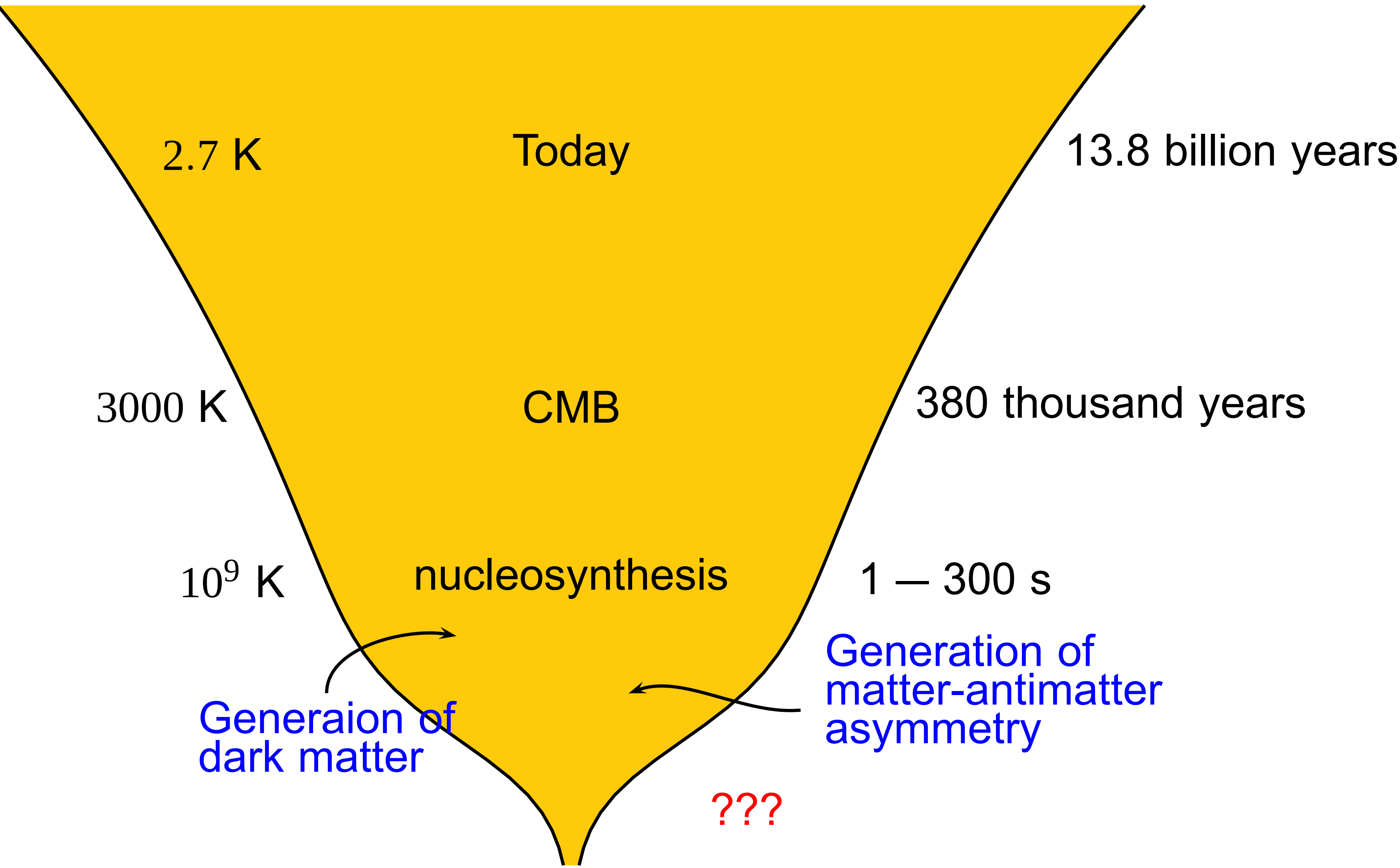
- Only very basic things are known for the time being.

## Good chance for future

- Detection of $B$-mode (partity odd) of CMB polarization $\Longrightarrow$ effect of primordial gravity waves $\Longrightarrow$ simple inflation
  - Together with scalar and tensor tilts $\Longrightarrow$ properties of inflaton
- Non-trivial correlation properties of density perturbations (non-Gaussianity) $\Longrightarrow$ contrived inflation, or something entirely different.
  - Shape of non-Gaussianity $\Longrightarrow$ choice between various alternatives
- Statistical anisotropy $\Longrightarrow$ anisotropic pre-hot epoch.
  - Shape of statistical anisotropy $\Longrightarrow$ specific anisotropic model

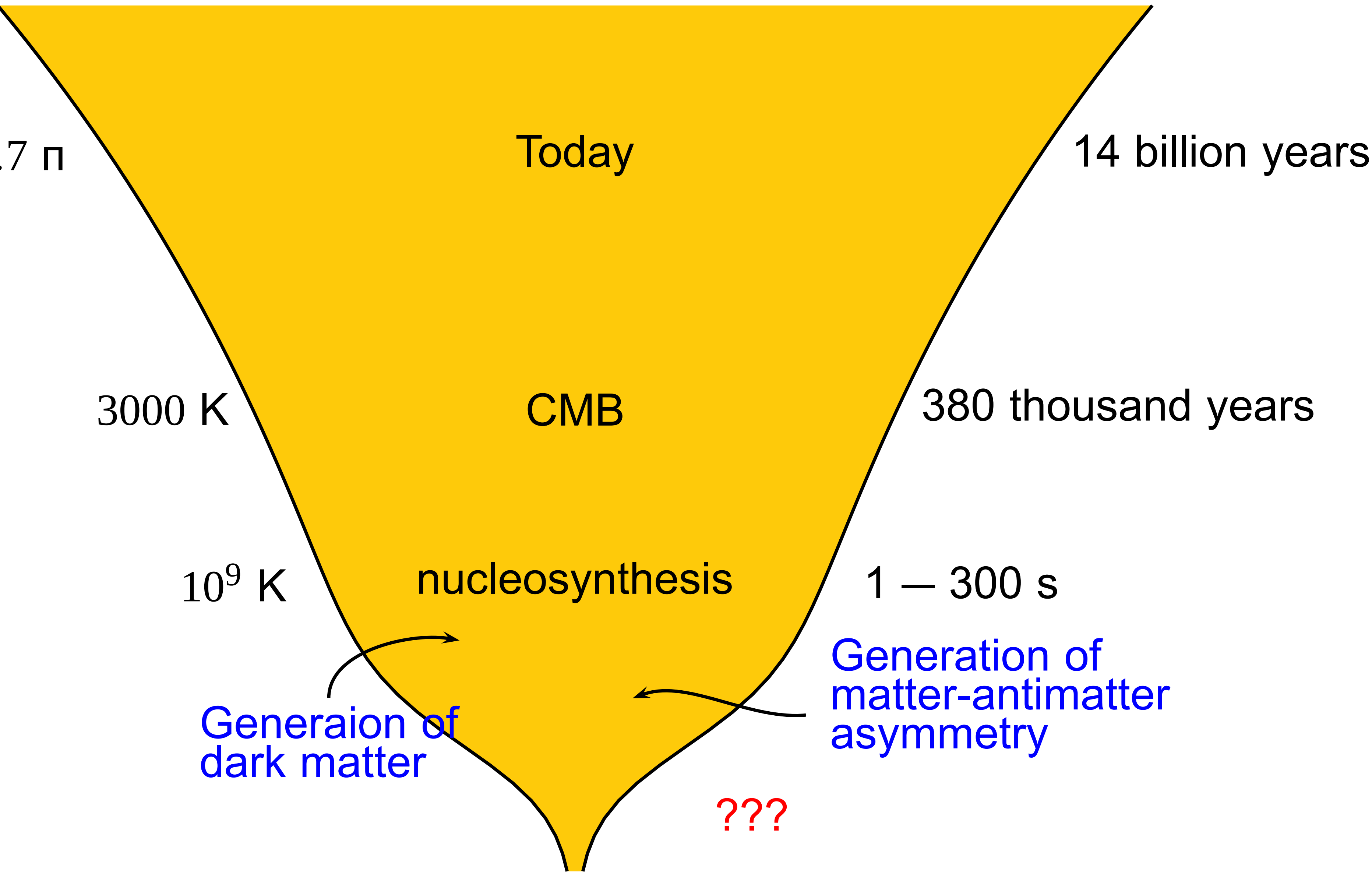
Good chance to learn
what preceeded the hot Big Bang epoch

Barring the possibility that Nature is dull



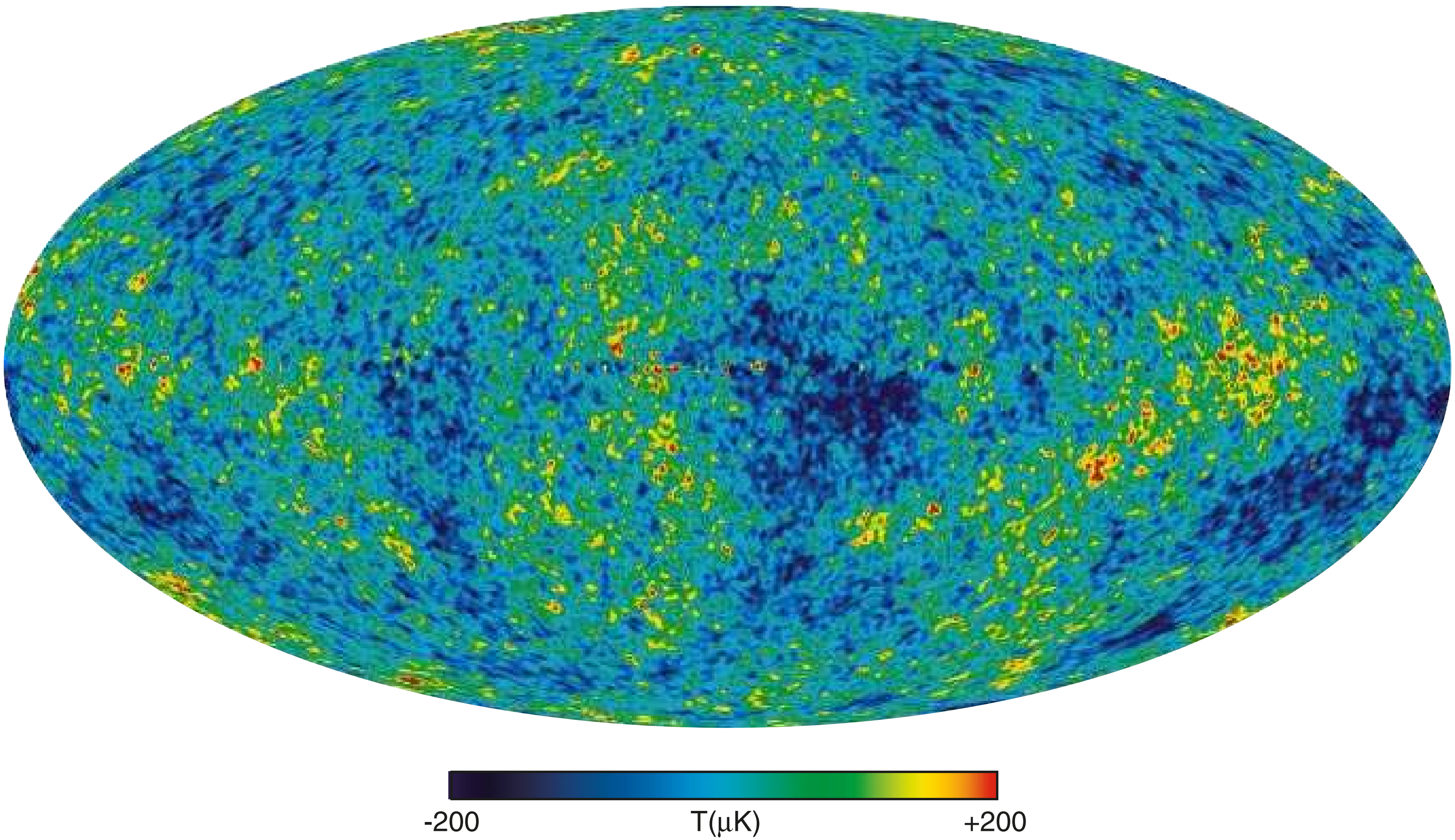
2.7 K
Today
13.8 billion years
3000 K
CMB
380 thousand years
$10^9$ K
nucleosynthesis
1 — 300 s
Generation of matter-antimatter asymmetry
Generaion of dark matter
???

# Backup slides


Today
14 billion years
3000 K
CMB
380 thousand years
$10^9$ K
nucleosynthesis
1 — 300 s
Generation of matter-antimatter asymmetry
Generaion of dark matter
???

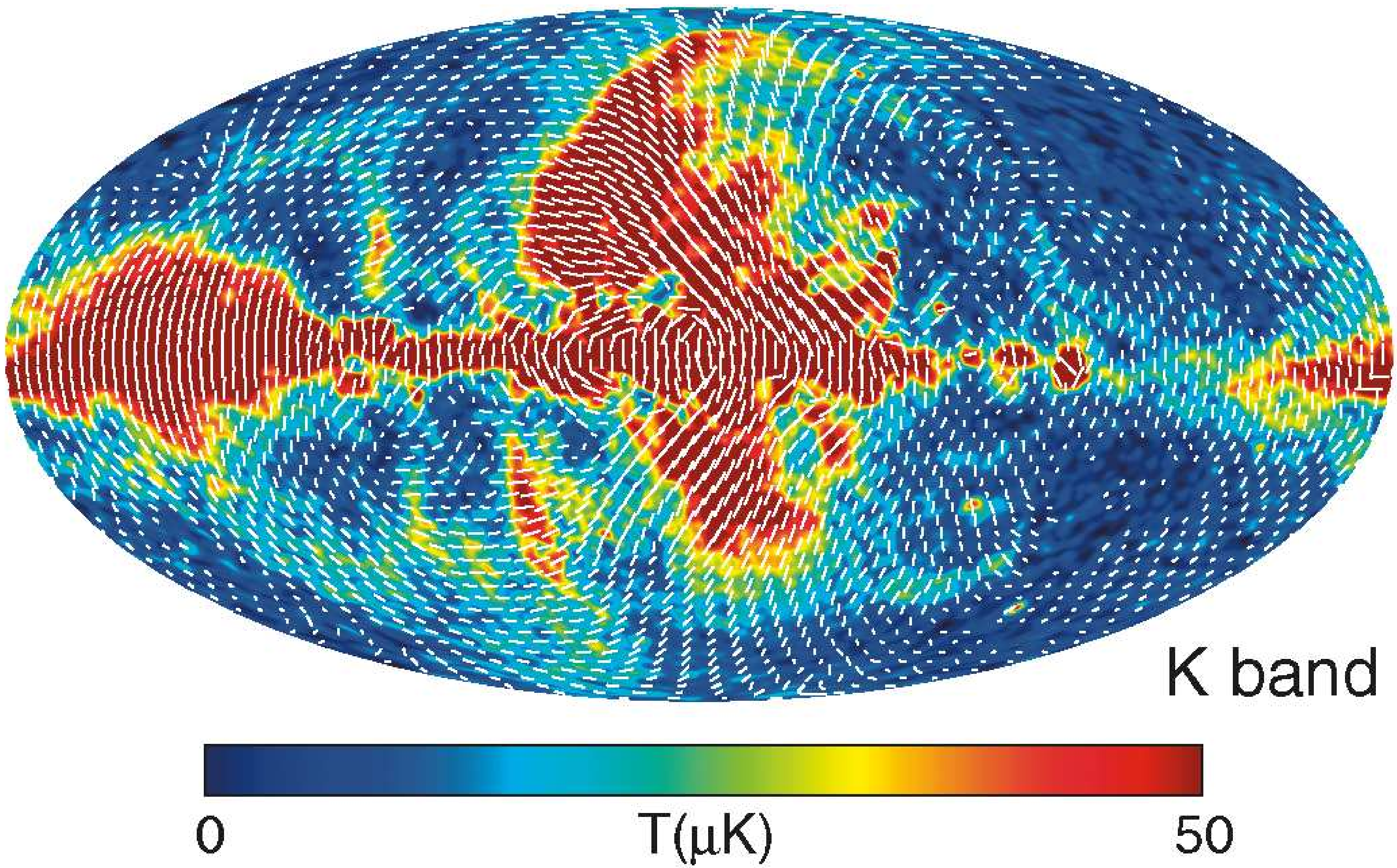
# CMB temperature anisotropy

$$T = 2.725^{\circ}K, \quad \frac{\delta T}{T} \sim 10^{-4} - 10^{-5}$$

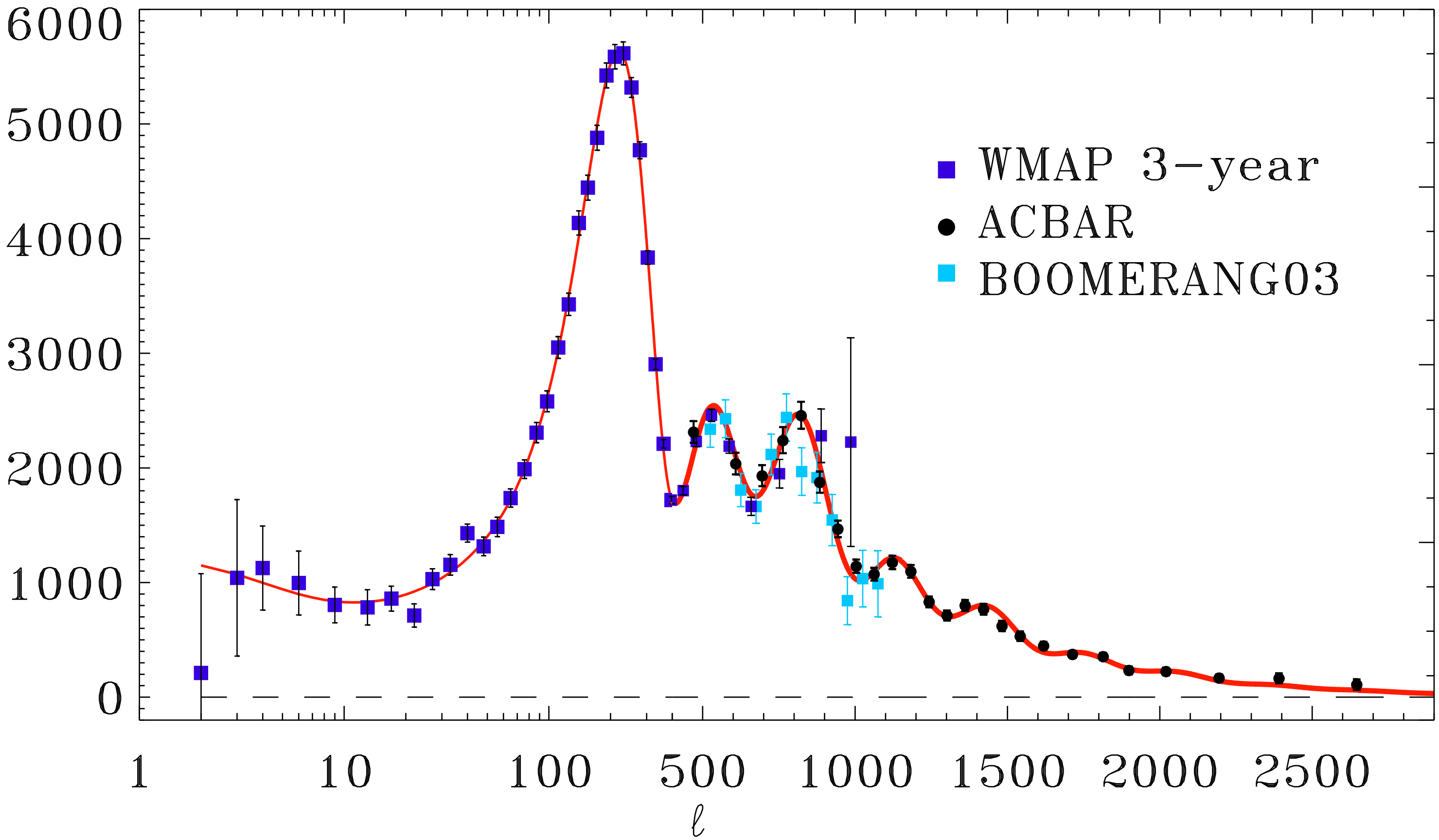



WMAP

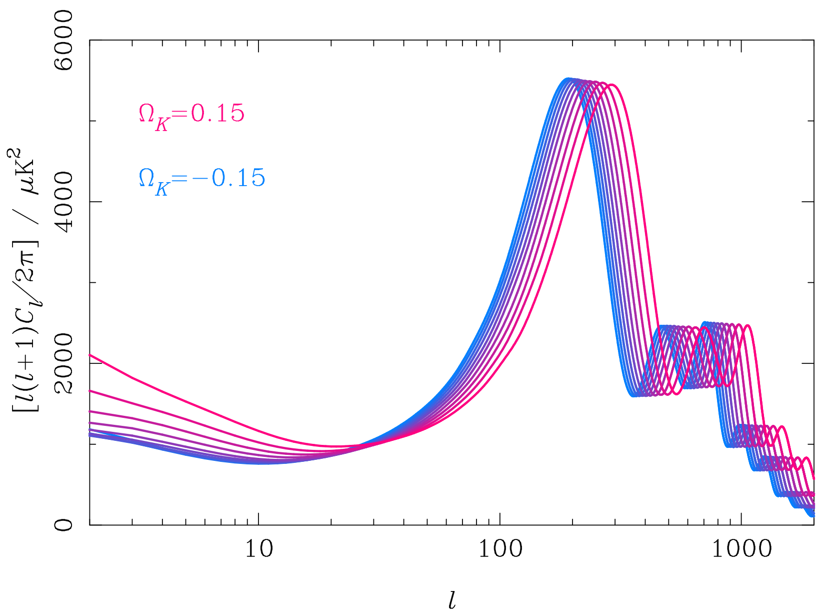
# CMB polarization map

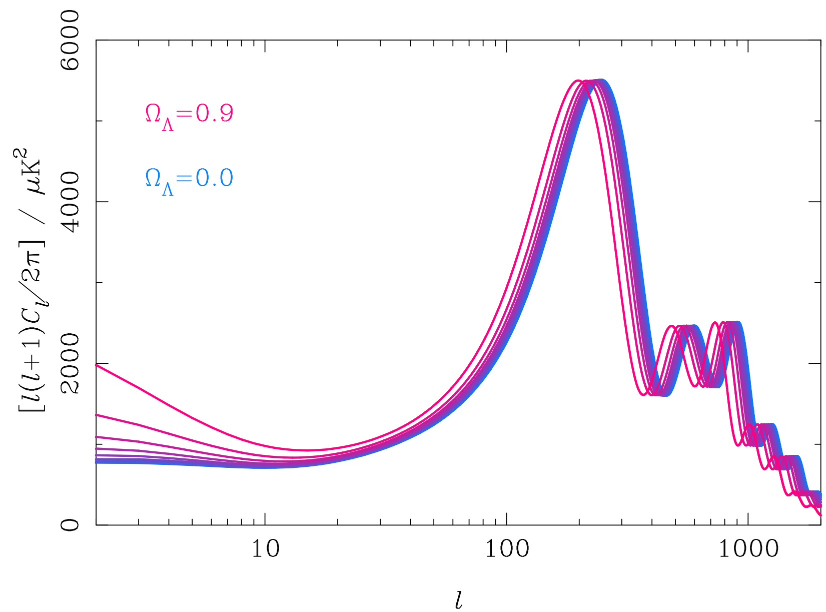


# CMB anisotropy spectrum

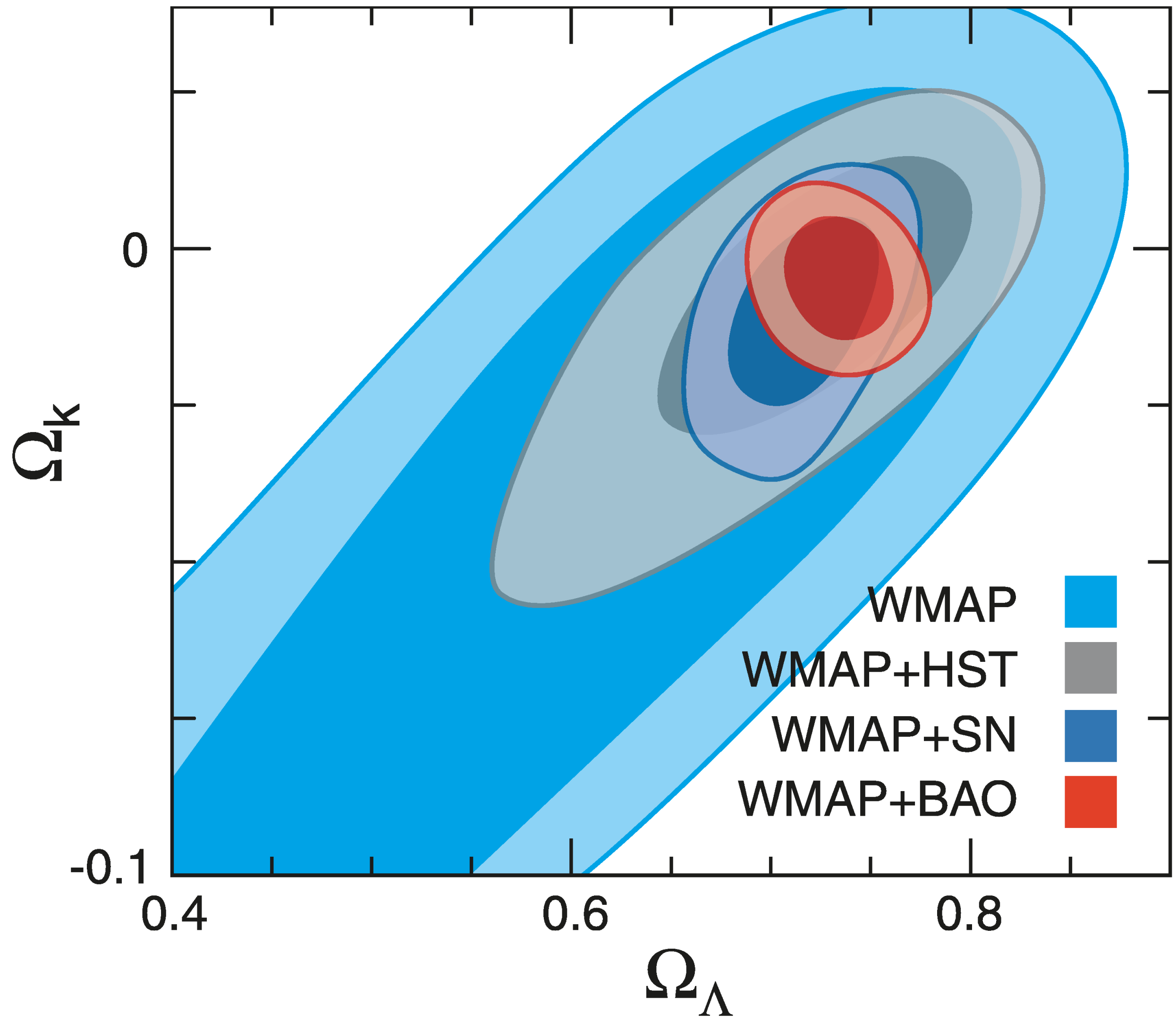


# Effect of curvature (left) and Λ

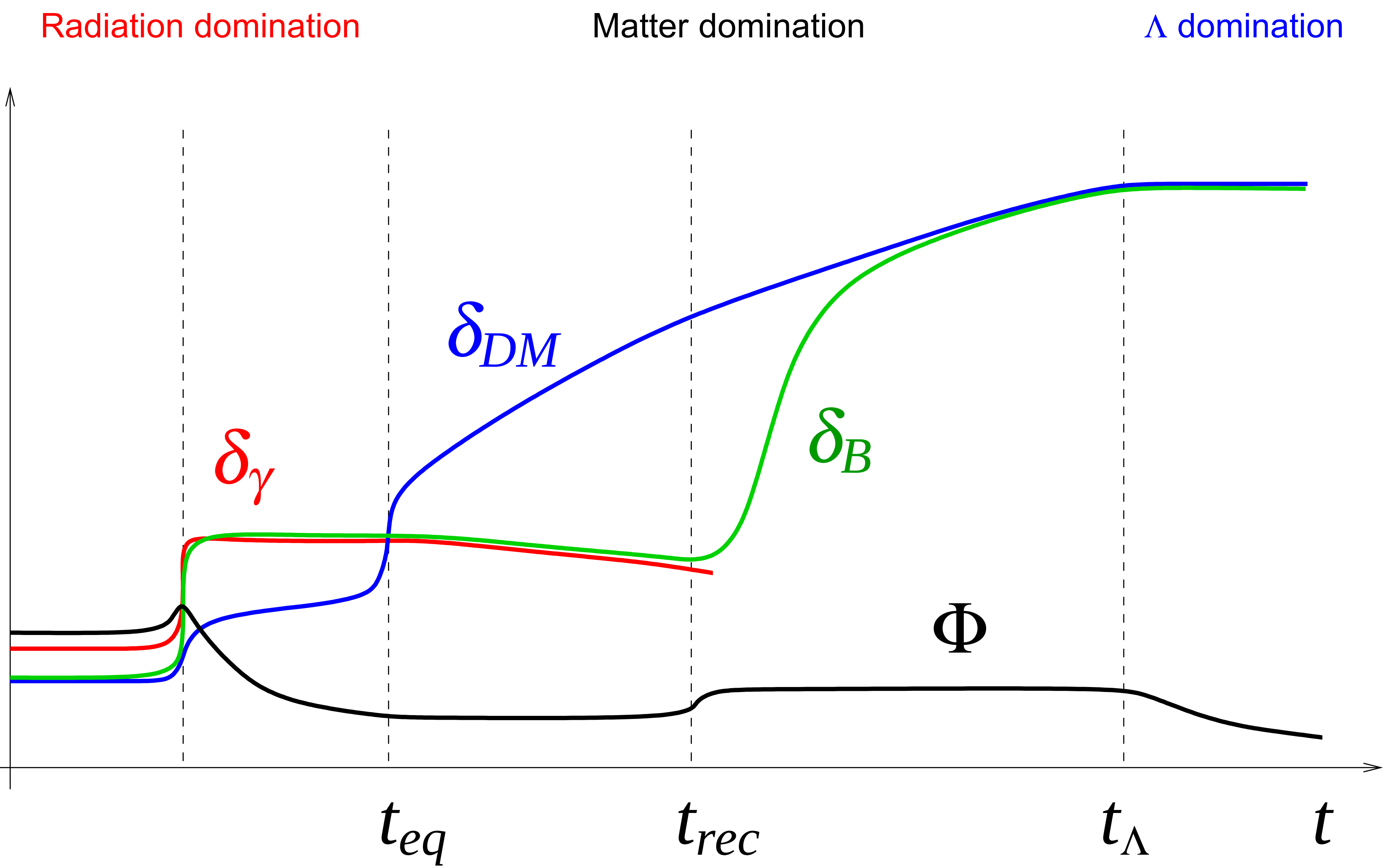


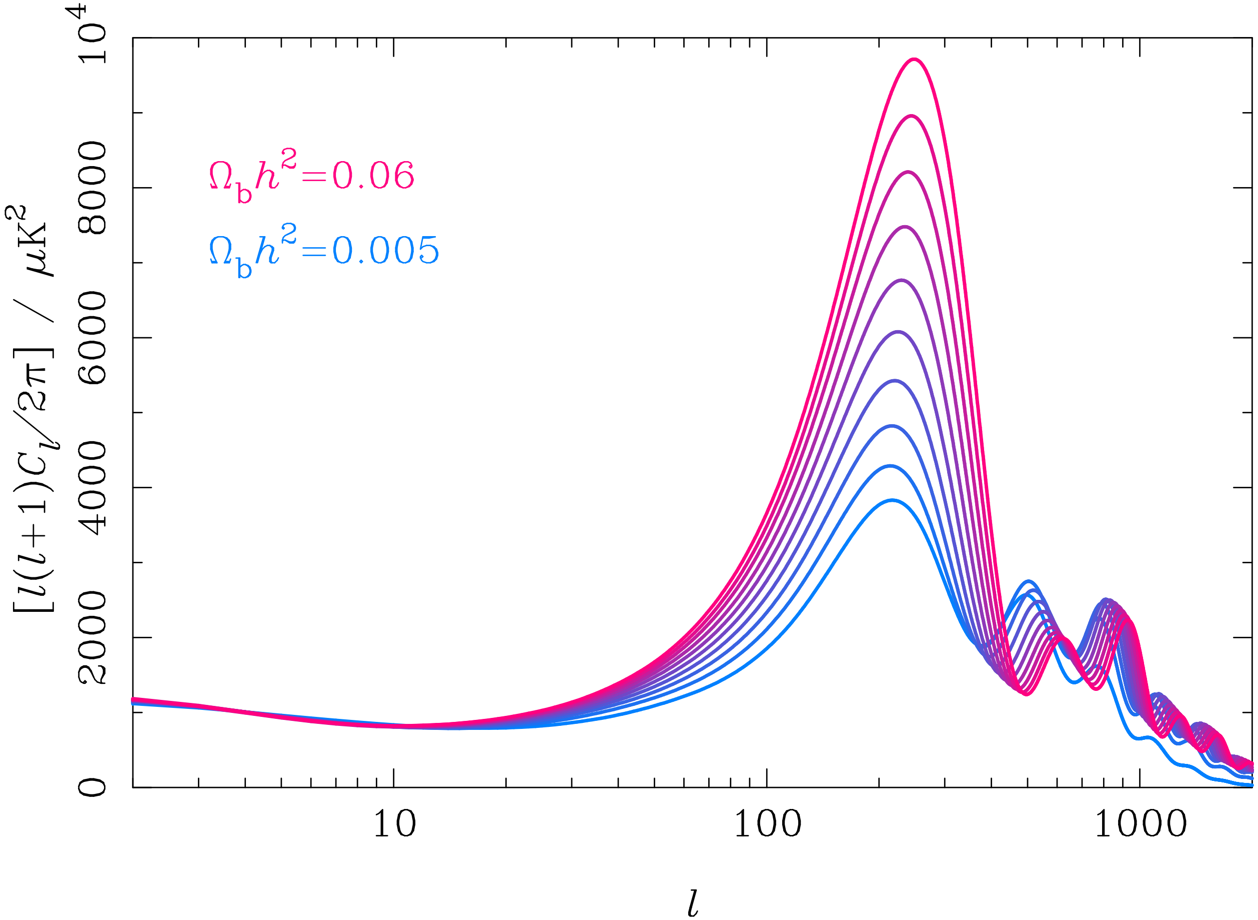
# Allowed curvature and Λ



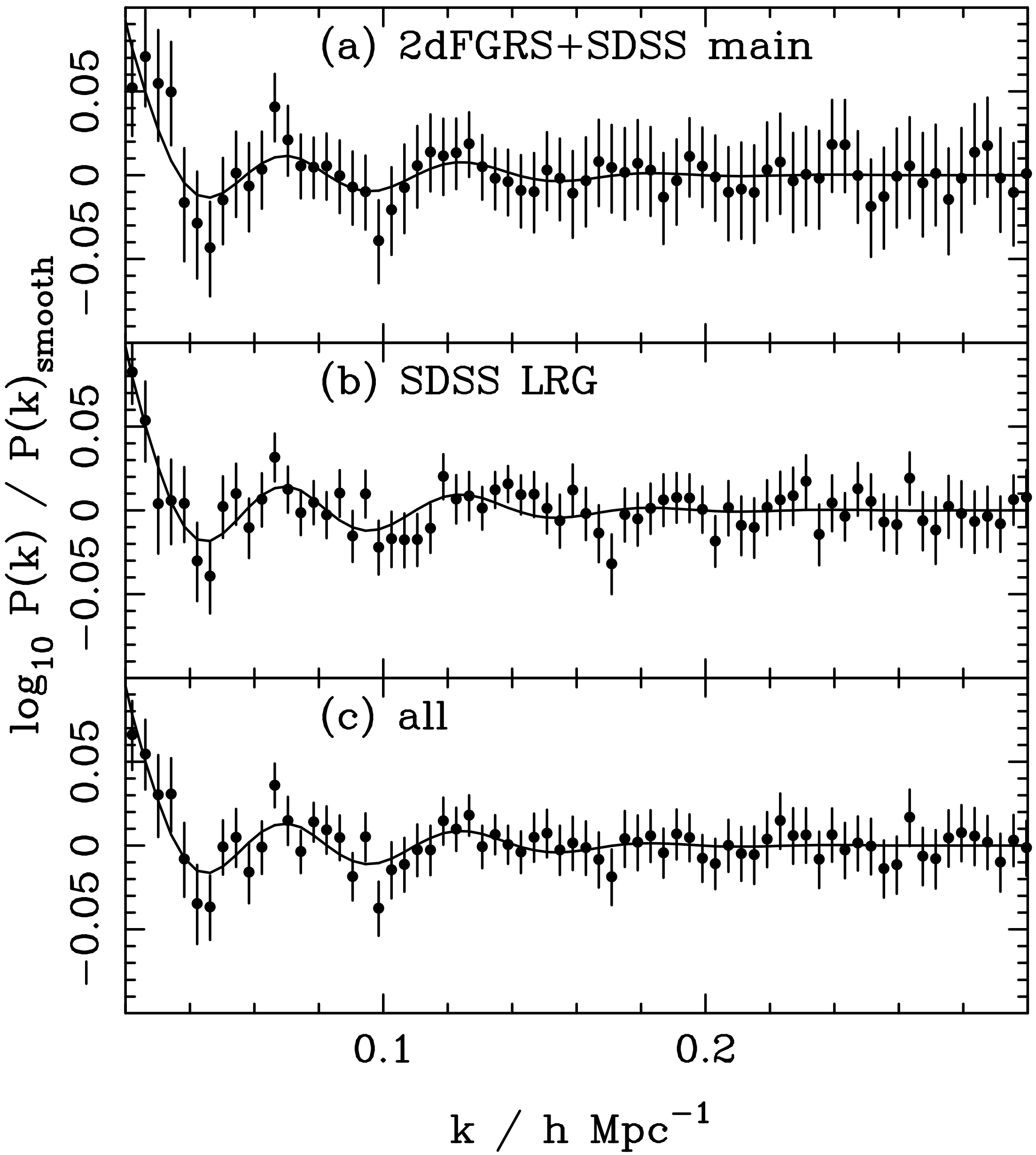
# Growth of perturbations (linear regime)



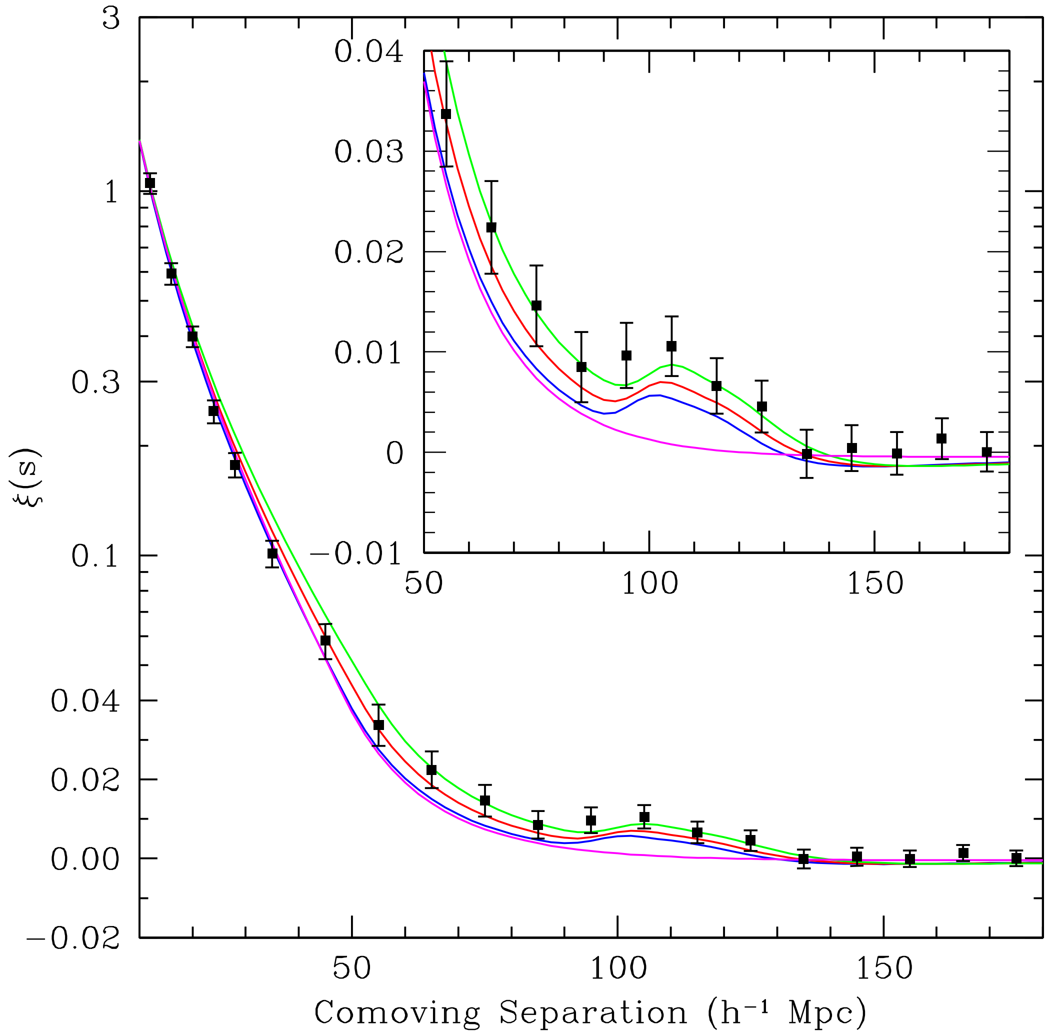
# Effect of baryons



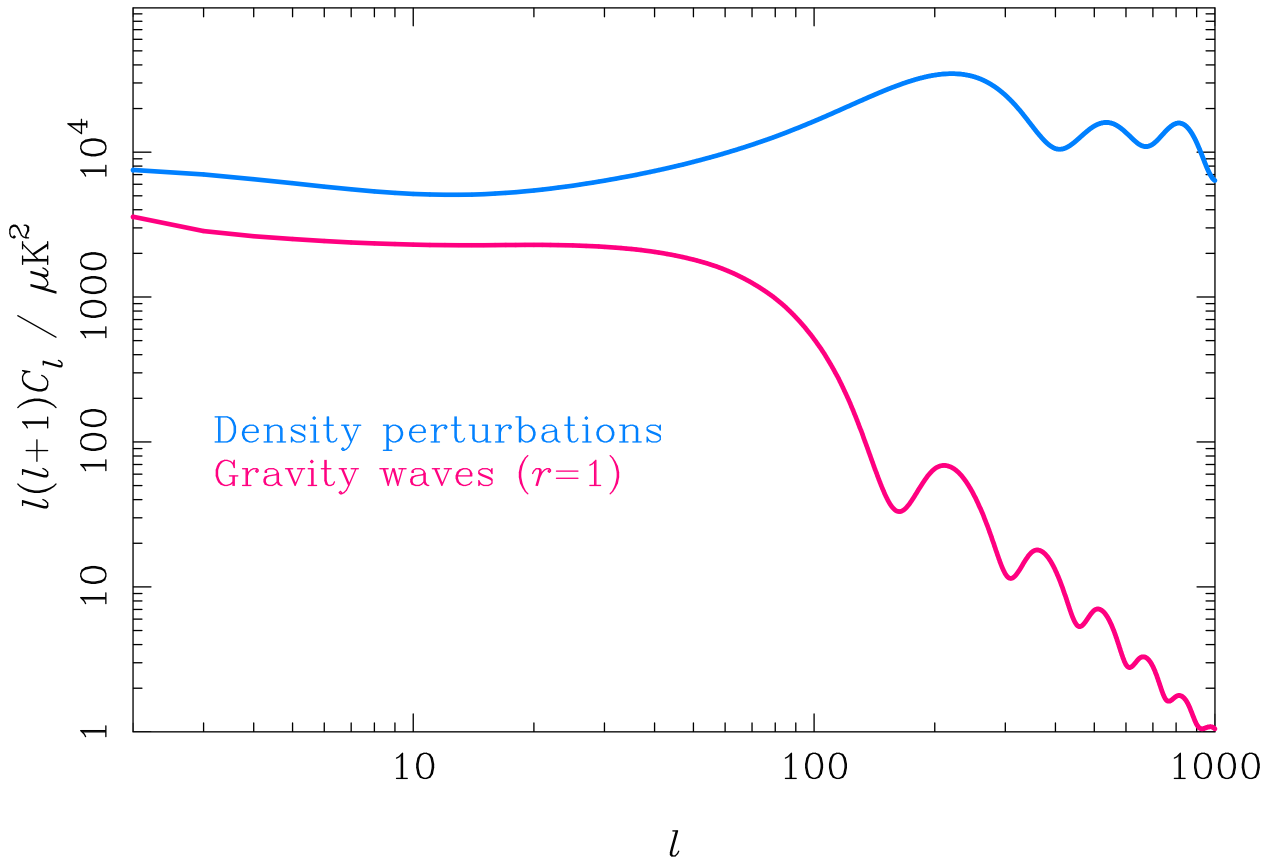
# BAO in power spectrum



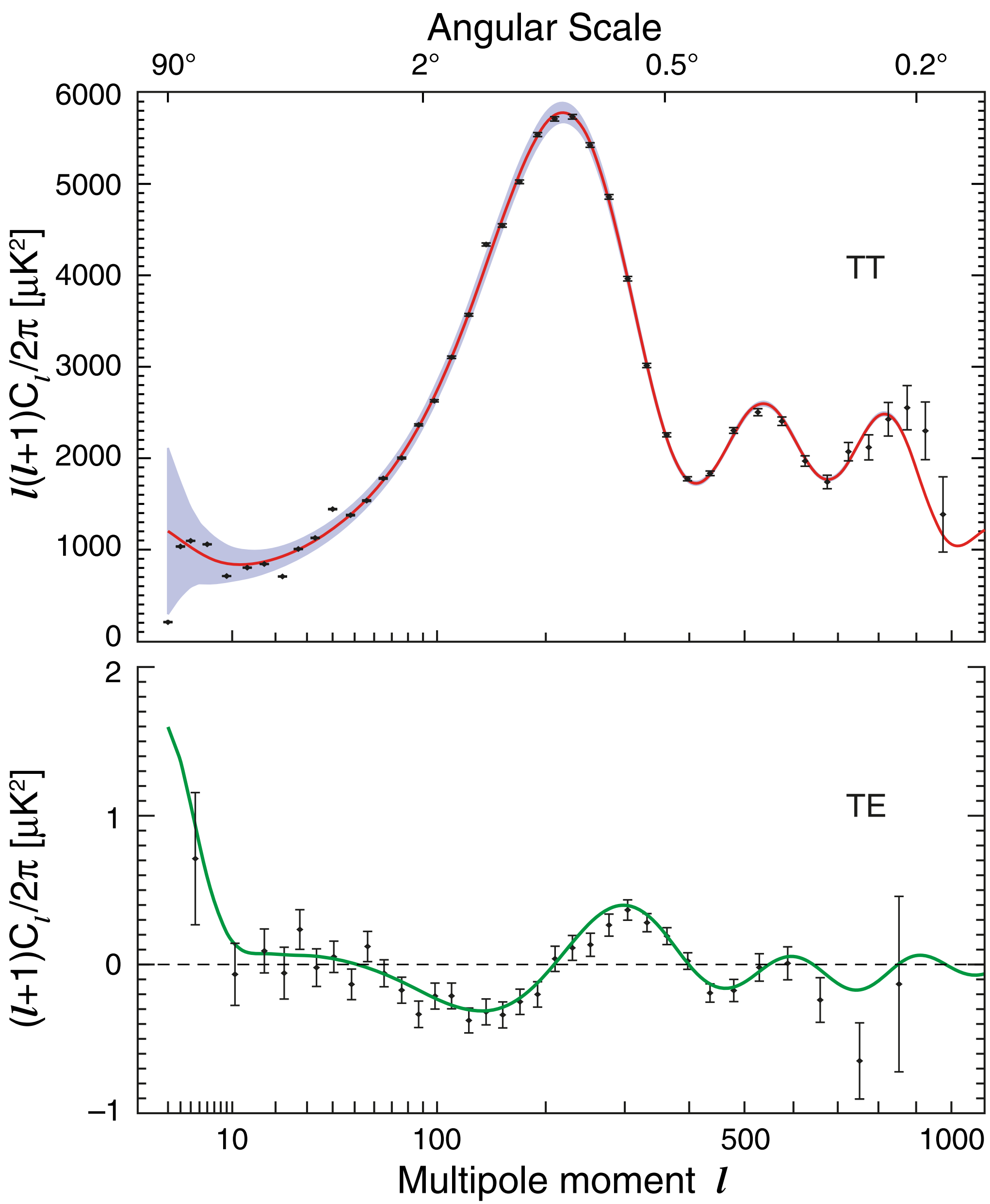
# BAO in correlation function



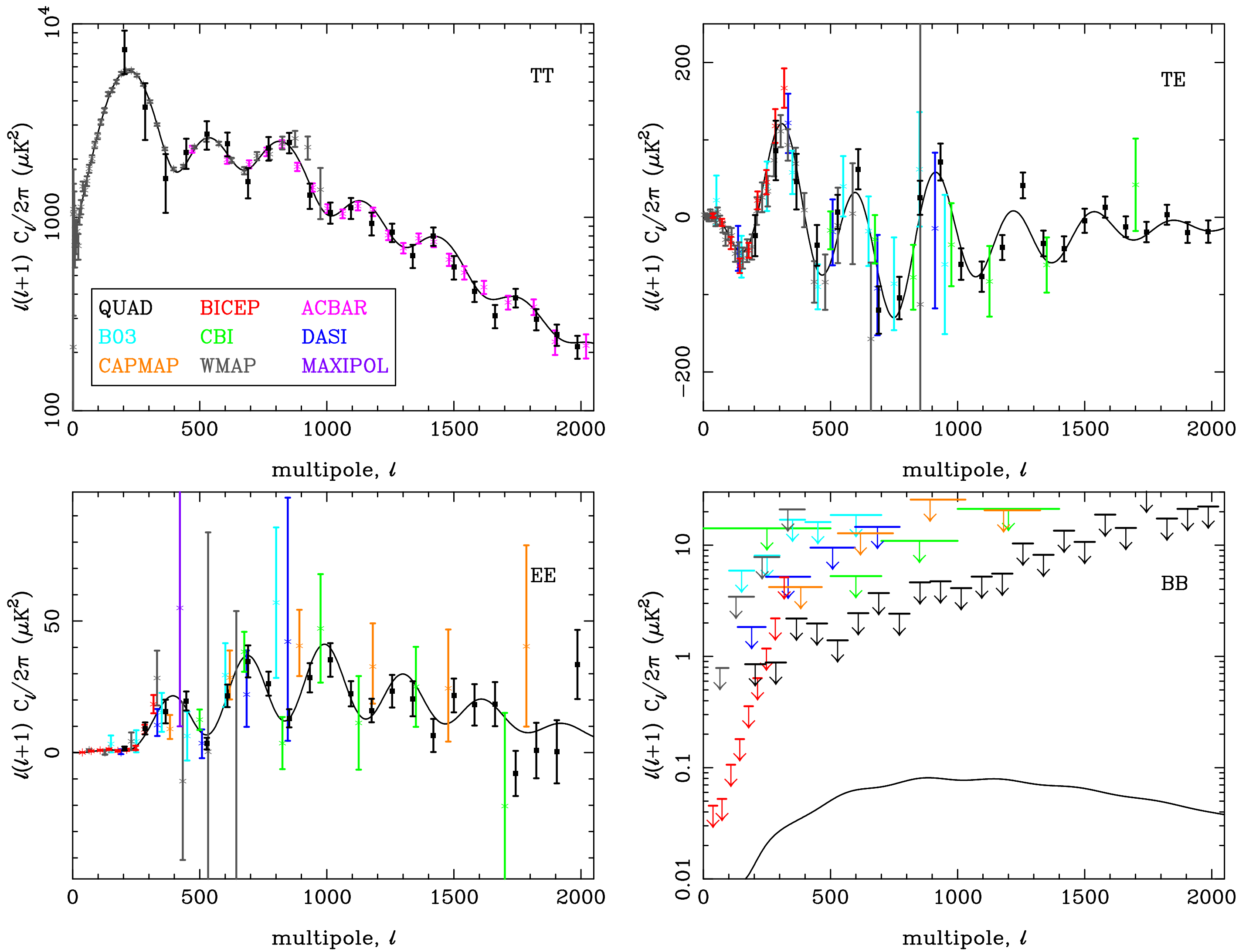
# Effect of gravity waves



# CMB temperature and polarization



# CMB temperature and polarization



# Effect of gravity waves on polarization (right)